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The binary Be star δ Sco at high spectral and spatial resolution[★]

Disk geometry and kinematics before the 2011 periastron

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ABSTRACT

Context. Classical Be stars are hot non-supergiant stars surrounded by a gaseous circumstellar disk that is responsible for the observed IR-excess and emission lines. The influence of binarity on these phenomena remains controversial.

Aims. δ Sco is a binary system whose primary suddenly began to exhibit the Be phenomenon at the last periastron in 2000. We want to constrain the geometry and kinematics of its circumstellar environment.

Methods. We observed the star between 2007 and 2010 using spectrally-resolved interferometry with the VLTI/AMBER and CHARA/VEGA instruments.

Results. We found orbital elements that are compatible with previous estimates. The next periastron should take place around July 5, 2011 (± 4 days). We resolved the circumstellar disk in the $H\alpha$ (FWHM = 4.8 ± 1.5 mas), $Br\gamma$ (FWHM = 2.9 ± 0.5 mas), and the $2.06\mu\text{m}$ He I (FWHM = 2.4 ± 0.3 mas) lines as well as in the K band continuum (FWHM ≈ 2.4 mas). The disk kinematics are dominated by the rotation, with a disk expansion velocity on the order of 0.2 km s^{-1} . The rotation law within the disk is compatible with Keplerian rotation.

Conclusions. As the star probably rotates at about 70% of its critical velocity the ejection of matter doesn't seem to be dominated by rotation. However, the disk geometry and kinematics are similar to that of the previously studied quasi-critically rotating Be stars, namely α Ara, ψ Per and 48 Per.

Key words. Techniques: high angular resolution – Techniques: interferometric – Stars: emission-line, Be – Stars: winds, outflows – Stars: individual (δ Sco) – Stars: circumstellar matter

1. Introduction

Classical Be stars are close-to-main-sequence hot stars that show or have shown infrared-excess and emission lines in their spectra. Such features stem from a dense gaseous circumstellar environment. However, the presence of a diluted wind with a terminal velocity of several hundreds of km s^{-1} have also been deduced from ultraviolet spectroscopy (Marlbrough & Peters 1986). Consequently, taking into account the fact that they are also fast rotators, a generally accepted view of these objects is that their circumstellar environment consists of two distinct regions: a dense equatorial disk dominated by rotation where most of the infrared radiation and emission lines are produced and a more diluted polar wind responsible for the highly broadened ultraviolet lines.

There is general agreement on the importance of stellar rotation on the ejection of matter and the break of the spherical symmetry of the circumstellar environment. However, the question

of whether or not additional physical processes are needed to produce ejection remains uncertain. Taking into account effects of gravitational darkening, Frémat et al. (2005) derived a mean value for the average rotational velocity of Be stars on the order of 88 % of their critical velocity (V_c) thus insufficient to fully explain the ejection of matter. In a statistical study of the rotation velocity of 462 Be stars Cranmer (2005) found that early type Be stars (O7-B3) exhibit a roughly uniform spread of intrinsic rotation speed extending from 40 % up to 100 % V_{crit} , whereas late types (B3-A0) are all quasi-critical rotators. This is a strong clue that rotation is essential in the formation of circumstellar envelopes for late type Be stars, whereas other physical mechanisms such as radiative pressure, pulsations or magnetism may dominate the ejection of matter for some of the earlier spectral type Be stars.

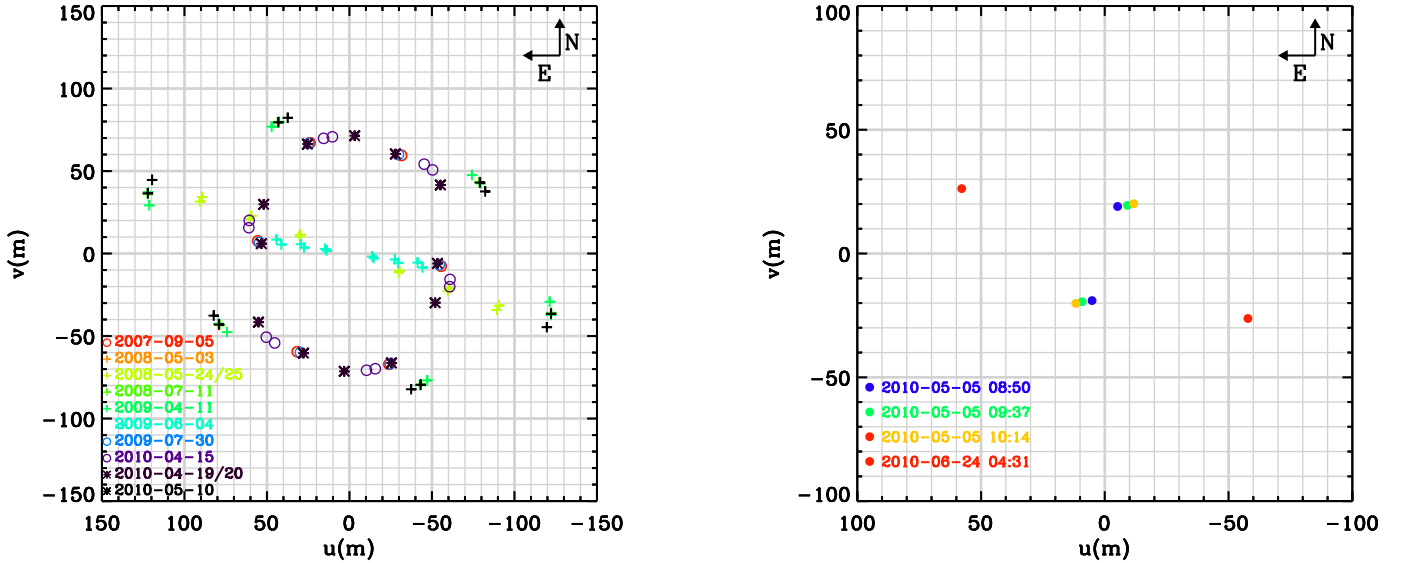
First VLTI/AMBER observations of Be stars α Ara (Meilland et al. 2007a) and κ CMa (Meilland et al. 2007b) have shown evidence of different geometries and kinematics, reinforcing the hypothesis of the heterogeneity of this group of stars in terms of mass-ejection processes (Stee & Meilland 2009). Moreover, recent discoveries of a companion around

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[★] Based on observations made with ESO Telescopes at Paranal under programs 383.D-0210, 385.D-0275, and CHARA/VEGA observations

Table 1. VLTI/AMBER and CHARA/VEGA observing logs for δ Sco.

VLTI/AMBER							
Observing Time Start (UTC)	Telescopes	Base Length (m)	Position Angle ($^{\circ}$)	Mode	Exposure/Frame (s)	Seeing ($''$)	Calibrators (HD)
2007-09-05 23:52	D0-H0-G1	67/56/71	-28/82/20	MR-K-F	1.00	0.67	146791
2008-05-03 06:12	A0-K0-G1	90/90/127	-151/-61/-106	LR-HK	0.05	0.79	139663
2008-05-24 04:28	A0-D0-H0	64/32/96	-109/-109/-109	LR-HK	0.05	0.99	139663, 132150
2008-05-25 04:05	A0-D0-H0	64/32/96	-110/-110/-110	LR-HK	0.05	0.81	139663
2008-07-11 01:48	A0-K0-G1	90/90/127	-151/-61/-106	LR-HK	0.05	0.97	139663, 166295
2009-04-11 07:40	A0-K0-G1	90/90/128	-151/-62/-106	LR-HK-F	0.05	0.63	139663
2009-04-11 08:19	A0-K0-G1	90/88/125	-148/-57/-103	LR-HK-F	0.05	0.82	139663
2009-06-04 05:20	E0-G0-H0	30/15/45	-100/-100/-100	LR-HK-F	0.05	0.74	139663
2009-06-04 05:20	E0-G0-H0	28/14/42	-97/-97/-97	LR-HK-F	0.05	0.94	139663
2009-07-30 02:27	D0-H0-G1	67/55/71	-26/83/20	MR-K-F	1.00	0.76	139663
2010-04-15 07:12	D0-H0-G1	64/71/71	72/-172/135	MR-K-F	0.50	0.82	139663
2010-04-15 07:55	D0-H0-G1	63/71/71	76/-167/140	MR-K-F	0.50	0.94	139663
2010-04-19 05:14	D0-H0-G1	60/71/69	60/177/127	HR-K-F (2.17 μ m)	6.00	0.85	139663
2010-04-20 09:14	D0-H0-G1	54/71/66	84/-159/155	HR-K-F (2.06 μ m)	6.00	0.60	139663
2010-05-10 05:11	D0-H0-G1	128/90/90	-110/115/-156	LR-K	0.05	0.83	139663
2010-05-10 05:52	D0-H0-G1	126/90/90	-107/119/-90	LR-K	0.05	0.95	139663
CHARA/VEGA							
Observing Time Start (UTC)	Telescopes	Length (m)	Position Angle P($^{\circ}$)	Mode	Detector	r ₀ (cm)	Calibrators (HD)
2010-05-05 08:50	S1-S2	20	-16	MR	Red	17	144470
2010-05-05 09:37	S1-S2	21	-25	MR	Red	18	144470
2010-05-05 10:14	S1-S2	23	-30	MR	Red	17	144470
2010-06-24 04:31	E1-E2	63	-111	MR	Red	9	144470

**Fig. 1.** (u,v) coverage for the complete VLTI/AMBER (left) and CHARA/VEGA (right) 2007-2010 dataset.

Achernar (α Eri, Kervella et al. 2008) and δ Cen (Meilland et al. 2008) may indicate that the putative effect of binarity on the Be phenomenon might have been underestimated. In the case of Achernar, Kanaan et al. (2008) successfully modeled spectroscopic variations as a brief equatorial outburst propagating into the circumstellar environment. Such phenomena probably originate from the close encounter between the central star and its companion.

In this context, a detailed study of δ Sco (HD 143275, HIP 78401) brings new perspectives to the understanding of the Be

phenomenon. This bright southern object has long been studied and first evidence of its multiplicity was reported by Innes (1901) using the lunar occultation technique. However, this work was forgotten for a long time, and the binary nature of δ Sco was rediscovered with three different techniques in 1974: by speckle-interferometry (Labeyrie et al. 1974), lunar occultation (Dunham 1974), and Intensity interferometry (Hanbury Brown et al. 1974). Using various interferometric measurements Bedding (1993) deduced an orbit of 10.6 years. However, the δ Sco system did not show clear evidence of the Be phenomenon

until the last periastron in June 2000. At this epoch, Otero (2001) found a 0.4 mag brightening of the object. Simultaneous spectroscopic observations published in Fabregat et al. (2000) showed evidence of strong H α emission lines.

In this work, we present near-infrared VLTI/AMBER (Petrov et al. 2007) and visible CHARA/VEGA (Mourard et al. 2009) interferometric measurements of δ Sco. The paper is organized as follows. In Sect. 2 we present the observations and the data reduction process. In Sect. 3 modeling of the data in the continuum allows us to constrain the binary orbit and physical parameters. In Sect. 4 we model the emission lines and constrain the envelope geometry and kinematics. Finally a short discussion in Sect 5 is followed by conclusions in Sect. 6.

2. Observation and data reduction

2.1. VLTI/AMBER

We initiated an interferometric follow-up of δ Sco soon after the opening of the VLTI/AMBER instrument to the scientific community. The star was then regularly observed between 2007 and 2010. The corresponding observing log is given in Table 1. Since this target is bright enough, i.e. $m_H \approx m_K \approx 2.4$, all observations were carried out using the 1.6 m auxiliary telescopes (AT). Low Resolution ($R = 30$), Medium Resolution ($R = 1500$), and High Resolution ($R = 12000$) spectral modes were used during the campaigns. The VLTI fringe tracker FINITO was used for all observations except those in 2008.

The data were reduced using the AMBER data reduction software *amdlib*, version 2.2. (Tatulli et al. 2007). The average raw complex visibility and closure phase was determined using the standard method, keeping the 20% of the frames with the higher SNR ratio. The interferometric calibration was then done using custom scripts described in Millour et al. (2008)¹. The (u,v) coverage for all observations is plotted in Fig 1.

Oscillations due to the binarity of δ Sco are seen in all LR and MR observations except those of 2008-05-03. Moreover three emission lines (Br γ , He I 2.06 μ m and Br δ) Fig 2 are clearly visible in the MR observations as shown in Fig. 2.

2.2. CHARA/VEGA

δ Sco ($m_V \approx 2.3$) was observed with the CHARA interferometer (ten Brummelaar et al. 2005) using the newly available VEGA instrument (Mourard et al. 2009). This 2-4 telescope visible beam-combiner has currently the highest spatial resolution ($\theta_{\min} = 0.3$ mas at 0.4 μ m with the 330m baseline) and highest spectral resolution for an optical/IR interferometer (up to $R = 30000$). Observations were carried out in May and June 2010 using the MR mode centered on the H α line (656.3nm). We obtained four measurements, three with the S1-S2 baseline (≈ 30 m), and one with the W1-W2 baseline (≈ 70 m). The log of the observations including the star used as an interferometric calibrator are shown in Table 1.

The data were reduced using the standard CHARA/VEGA reduction software developed by the VEGA team and described in Mourard et al. (2009). This software offers two data reduction modes, the first based on spectral density analysis and the second on the cross-spectral scheme.

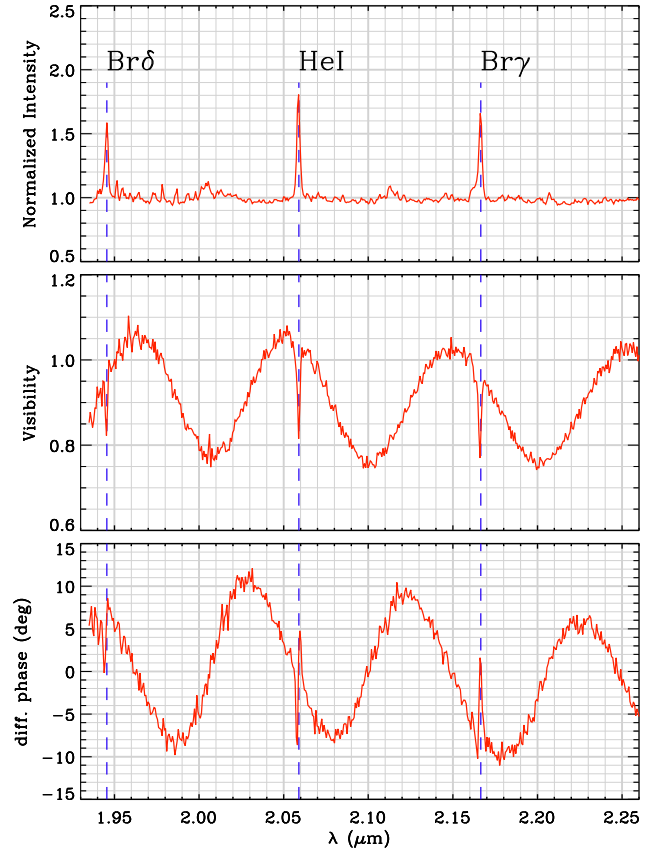


Fig. 2. Example of a spectrally resolved VLTI/AMBER measurement for one baseline in MR mode. The oscillations of the visibility and phase plotted as the function of the wavelength are due to the binarity of the object. This figure also exhibit effects of the δ Sco circumstellar environment geometry and kinematics in the Br γ , He I and Br δ emission lines.

3. The binary system

3.1. Calculating the binary separation

We used the LITpro² model fitting software for optical/infrared interferometric data developed by the Jean-Marie Mariotti Center (JMMC) to analyze our data. It is based on the Levenberg-Marquardt algorithm, that allows from a set of initial values of the model parameters to converge to the closet χ^2 local minimum. It also include tools to facilitate the search for the global minimum. **LITpro calculates an error on the fitted parameters based on the χ^2 value at the minimum. It uses data error estimates based on the OI FITS format, which does not include error correlation estimates. Therefore, in some cases, LITpro can provide underestimated errors on the parameters.**

Since the aim of this first study is to refine the previous determination of the binary orbit, we only used AMBER LR-K and MR-K data, where the binary signal is clearly visible. In our analysis, we also neglect signal coming from the K band emissions lines visible in Fig 2.

The primary is modeled as a uniform disk to take into account partial resolution of its circumstellar environment in the continuum, whereas the companion is modeled as a point like source. Consequently our model has only 4 free parameters: the

¹ The scripts are available to the community at the following web-page: <http://www.mpifr-bonn.mpg.de/staff/fmillour>

² LITpro software available at <http://www.jmmc.fr/litpro>

Table 2. 2007-2010 evolution of the δ Sco binary system determined from our VLTI/AMBER observations.

Date	Model parameters				χ_r^2	Separation in polar coordinates	
	x (mas)	y (mas)	F_1	D_1 (mas)		sep (mas)	θ (deg)
2007-09-05	-3.9±0.3	-168.9±0.3	0.96±0.06	1.7±0.4	2.8	168.9	1.3
2008-05-03	-	-	1.00	1.7±0.6	10	-	-
2008-05-24/25	-7.8±0.4	-163.5±1.0	0.96±0.01	1.4±0.1	6.1	163.7	2.7
2008-07-11	-6.5±0.4	-163.0±0.5	0.99±0.03	0.9±0.1	1.2	163.1	2.3
2009-04-11	-11.5±1.0	-156.8±1.0	0.98±0.02	2.0±0.1	3.6	157.2	4.2
2009-06-04	-12.7±0.5	-146.8±3.0	0.94±0.01	1.7±0.1	0.9	147.3	4.9
2009-07-30	-12.8±1.5	-135.3±2.5	0.93±0.01	2.7±0.1	2.3	135.9	5.4
2010-04-15	-27.7±0.1	-101.6±0.1	0.94±0.02	1.9±0.5	2.9	105.3	15.3
2010-05-10	-22.6±0.1	-96.3±0.1	0.95±0.1	1.4±0.1	2.4	98.9	13.2

binary separation in Cartesian coordinates (x and y) where x (resp. y) is counted positive toward the east (resp. north), the relative flux of the primary to the total flux (F_1), and its diameter (D_1).

The result of the model-fitting is presented in Table 2. The binary separation decreased from about 170 mas in 2007 to 100 mas in 2010. **The estimated error from LITpro, in average 0.8 mas on x and y , is reported in this table. The data points scatter around the fitted orbit (Sect. 3.2) is of the order of 7 mas. Therefore, our error estimates are very likely underestimated.** The K-band primary relative flux F_1 is varying from 92 to 100 % of the total flux. However, this variation is linked to the spatial frequency sampling since it depends on the baseline length and spectral resolution. The minimum values of F_1 are obtained for the MR observations. For the LR ones its value varies from $98 \pm 1\%$ for observations made with the longest triplet A0-K0-G1 to 94 % for the shorter triplet E0-G0-H0. By studying the measurements with the intermediate triplet A0-D0-H0, we clearly see a change in the oscillation amplitudes from the short baselines to the largest. Such a variation could have been a consequence of resolving the companion. However, since these oscillation amplitudes determined using observations made in MR and LR modes for the same baseline length are mutually inconsistent, this clearly indicates that it comes from a technical problem due to the sampling of oscillations. Finally, we can deduce that the amplitudes measured in MR mode are not biased by this spatial frequency sampling effect. Thus the real K-band relative flux of the primary is on the order of $93 \pm 1\%$.

According to almost all measurements, the primary component is partially resolved. Modeling it as a simple uniform disc we can derive a diameter of 1.6 ± 0.6 mas. Adopting the stellar radius from Carciofi et al. (2006), $7 R_\odot$, and considering a distance of 150 pc (van Leeuwen 2007), this corresponds to $25.7 \pm 9.6 R_\odot$, i.e. $3.67 R_\star$. This gives us a lower limit of the circumstellar disk extension in the K band for the case where it fully dominates the primary K band flux. If only half of the primary K band emission would originate from the disk, as for α Ara (Meilland et al. 2007a) or κ CMa (Meilland et al. 2007b), its extension would be on the order of 2.4 mas, i.e. $4.6 R_\star$.

3.2. The orbital elements

3.2.1. Previous work

Since the rediscovery of its binarity in 1974, δ Sco has been observed several times using speckle-interferometry, mainly published in McAlister & Harkopf (1988). A first attempt to constrain its orbit using this set of data as well as optical-interferometry data from the MAPPIT experiment was by

Bedding (1993). New speckle-interferometric measurements allowed Hartkopf et al. (1996) to derive more accurate and significantly different parameters. Using the same dataset complemented by radial velocity measurements close to the 2000 periastron, Miroshnichenko et al. (2003) refined the previous analysis. Tango et al. (2009) have used the same dataset but with a more consistent model-fitting algorithm from Pourbaix (1998) to estimate the orbit. Finally the most recent analysis of δ Sco orbit was by Tycner et al. (2011) using 96 measurements from taken with the Navy Prototype Optical Interferometer (NPOI). The resulting orbital elements determined in these five papers are presented in Table 3.

The periastron passage (T_0) and the eccentricity are extremely well constrained by the radial-velocity measurements and small uncertainties remain for the orbit angular parameters i , ω , and Ω . The main differences between the values of the orbital elements derived by two latest and most complete studies are the orbit period (P) and major-axis (a). One important parameter for the ongoing δ Sco observing campaigns is the exact date of the next periastron. According to Miroshnichenko et al. (2001) it should take place in 2011 on April 9, whereas Tango et al. (2009) predict that this will occur between May 30 and June 14 and Tycner et al. (2011) that the next periastron passage is expected to occur on UT 2011 July 6 ± 2 d.

3.2.2. This work

All our measurements are plotted in Fig 3. They are roughly compatible with the most recent orbits determined by Tango et al. (2009) and Tycner et al. (2011) that are over-plotted on this figure. Other measurements available from the *Fourth Catalog of Interferometric Measurements of Binary Stars*³ are also over-plotted.

In Cartesian coordinates the χ^2 for the separation measurements is given by:

$$\chi_{\text{sep}}^2(P, T_0, a, e, i, \omega, \Omega) = \sum_{i=1}^{N_{\text{sep}}} \left[\left(\frac{x_i - \hat{x}_i}{\sigma_{x,i}} \right)^2 + \left(\frac{y_i - \hat{y}_i}{\sigma_{y,i}} \right)^2 \right] \quad (1)$$

where x_i and y_i is the measured binary separation in Cartesian coordinates for the $N_{\text{sep}}=150$ measurements (i.e. 96 from Tycner et al 2011, 46 from the USNO database, and 8 from our measurements), $\sigma_{x,i}$ and $\sigma_{y,i}$, the uncertainties of the measurements, \hat{x}_i and \hat{y}_i the modeled separation at the corresponding epochs. χ_{sep}^2 is a function of the 7 orbital parameters: P , T_0 , a , e , i , ω , and Ω .

³ <http://ad.usno.navy.mil/wds/int4.html>

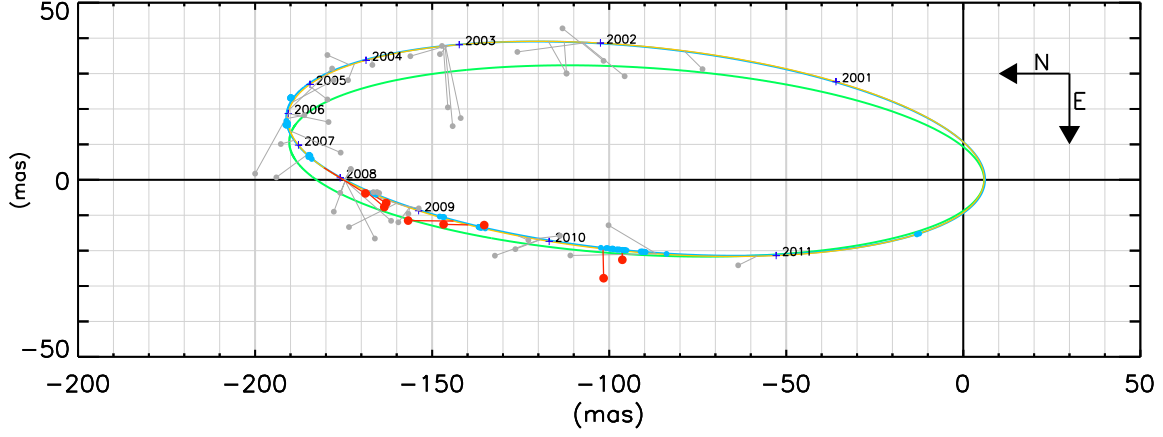


Fig. 3. The visual orbit of δ Sco. The gray circles are binary separations taken from the *Fourth Catalog of Interferometric Measurements of Binary Star*, the blue ones from Tycner et al. (2011) and the red ones are our new measurements derived from our VLT/AMBER data. The green and blue solid lines represent the orbits derived by Tango et al. (2009) and Tycner et al. (2011), respectively. Our best-fit orbit is plotted as an orange solid line.

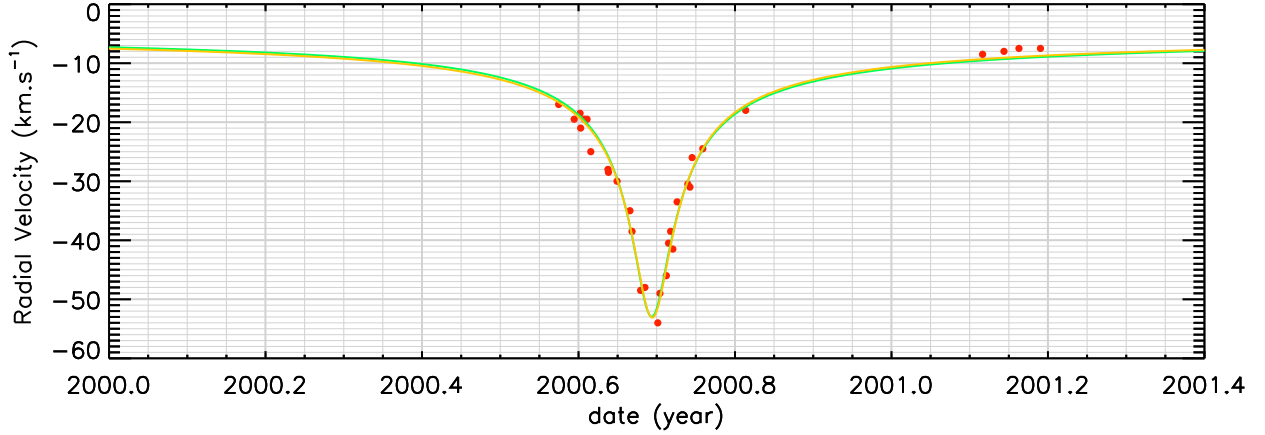


Fig. 4. δ Sco radial velocity around the last periastron. Data (red circles) are taken from Miroshnichenko (2001). The green and orange lines represent Tango et al. (2009) and our best-fit model, respectively.

As in Tango et al. (2009) we also use radial velocities from Miroshnichenko et al. (2001) plotted in Fig 4 in our model-fitting process. The χ^2 for these measurements is given by:

$$\chi_{v_{\text{rad}}}^2(T_0, P, e, \omega, V_0, K_A) = \sum_{i=1}^{N_{v_{\text{rad}}}} \left(\frac{v_{\text{rad}i} - \widehat{v_{\text{rad}i}}}{\sigma_{v_{\text{rad}i}}} \right)^2 \quad (2)$$

where $v_{\text{rad}, i}$ is the measured radial velocity for the $N_{v_{\text{rad}}} = 30$ measurements, $\sigma_{v_{\text{rad}, i}}$ the corresponding uncertainty on the measurements, and $\widehat{v_{\text{rad}i}}$ the modeled radial velocities. $\chi_{v_{\text{rad}}}^2$ is a function of the orbital parameters T_0 , P , e , and ω and of two additional parameters: the systemic radial velocity of the system V_0 , and the amplitude of the velocity variations K_A .

We simultaneously fit the visual separation measurements and the radial velocities by minimizing the function:

$$F(P, T_0, a, e, i, \omega, \Omega, V_0, K_A) = \chi_{\text{sep}}^2 + \chi_{v_{\text{rad}}}^2 \quad (3)$$

To minimize F we used the *downhill simplex method* by Nelder & Mead (1965). The starting parameters were set in four

different ways. In the three first fitting processes we used the orbital elements from Tango et al. (2009), Miroshnichenko et al. (2001), and Tycner et al. (2011). Finally, in the last fitting process we used 10^4 random starting positions. For each starting position the simplex algorithm was allowed to converge until the difference in χ^2 between two consecutive models is less than $\epsilon = 10^{-4}$.

Using these initial guesses the fit converged to the same solution. The best fit model parameters are shown in Table 3 and the corresponding visual orbit and radial velocity curve are overplotted on Figs 3 and 4, respectively. We found values that fully agrees with the latest estimates from Tycner et al. (2011). This is reasonable taking into account that their measurements dominates the sample both in term of number (96 out of 150) and uncertainties. Finally, we can conclude that the next periastron should take place around July 5, 2011 (± 4 days).

3.3. The binary system physical parameters

We can constrain the total mass of the δ Sco binary system using Kepler's third law and our estimation of the binary period

Table 3. The orbital elements of δ Sco.

Orbital el.	Bedding 1993	Hartkopf et al. 1996	Miroshnichenko et al. 2001	Tango et al. 2009	Tycner et al. 2011	This work
P (yr)	10.5	10.59 ± 0.075	10.58	10.74 ± 0.02	10.817 ± 0.005	10.811 ± 0.01
T_0 (yr)	1979.3	1979.41 ± 0.14	2000.693 ± 0.008	2000.69389 ± 0.00007	2000.6927 ± 0.0014	2000.6941 ± 0.003
a (mas)	110	106.7 ± 6.7	107	98.3 ± 1.2	99.1 ± 0.1	98.74 ± 0.07
e	0.82	0.92 ± 0.02	0.94 ± 0.01	0.9401 ± 0.0002	0.9380 ± 0.0007	0.9403 ± 0.0008
i ($^\circ$)	70	48.5 ± 6.6	38 ± 5	38 ± 6	32.9 ± 0.2	30.2 ± 0.7
ω ($^\circ$)	170	24 ± 13	-1 ± 5	1.9 ± 0.1	2.1 ± 1.1	0.7 ± 2.9
Ω ($^\circ$)	0	159.6 ± 7.6	175	175.2 ± 0.6	172.8 ± 0.9	174.0 ± 2.5
V_0 (km s $^{-1}$)			6	-6.72 ± 0.05	-7	-6.7 ± 0.2
K_A (km s $^{-1}$)				23.83 ± 0.05	?	23.9 ± 0.1

and semi major-axis. Since this last parameter is given in angular units (i.e. mas), the binary system distance is also needed for the mass calculation. Using the distance inferred from the Hipparcos parallax (Perryman et al. 1997), i.e. $d = 123$ pc, we derived a total mass of $M_{A+B} = 15.2 \pm 5 M_\odot$. On the other hand, using the distance deduced from van Leeuwen (2007) revised Hipparcos parallax, i.e. $d = 150$ pc, we obtained a total mass $M_{A+B} = 27.7 \pm 10 M_\odot$. The second estimation is in agreement with the total mass determined previously by Tango et al. (2009), i.e. $23 \pm 10 M_\odot$.

From physical modeling of the spectral energy distribution, Carciofi et al. (2006) estimated the mass of the primary to be on the order of $M_A = 14 M_\odot$. Assuming that both stars have the same age and are currently on the main sequence, and considering the flux ratio measured in the K band, i.e. $F_B/F_A = 0.07/0.93 = 0.075 \pm 0.012$, the secondary should also be a quite massive hot star. Thus, it seems that $M_{A+B} = 15.2 \pm 5 M_\odot$ is quite unrealistic. On the other hand $M_{A+B} = 27.7 \pm 10 M_\odot$ would implies that both components have approximatively the same mass.

Finally, Tango et al. (2009) give an estimation of the secondary mass of $M_B = 8 \pm 3.6 M_\odot$. From the mass ratio they calculated, i.e. $M_A/M_B = 1.957 \pm 0.011$ we can deduce the Luminosity ratio using Mass-Luminosity relation from Griffiths et al. (1988):

$$\frac{L_B}{L_A} = \left(\frac{M_B}{M_A} \right)^{3.51 \pm 0.14} \quad (4)$$

Thus, we can deduce that $L_B/L_A = 0.095 \pm 0.01$. This is compatible with our K band flux ratio measurements, i.e. $F_B/F_A = 0.075 \pm 0.012$. This gives an additional indication that the two components have similar spectral classes. Considering the measured flux ratio, this would imply that the secondary spectral class ranges between B2V and B4V.

4. The envelope geometry and kinematics

During our observing campaign, δ Sco was observed three times with AMBER in medium resolution (MR): in 2007, 2009, and 2010. All the observations were centered on the fist half of the K band (1.95 - $2.25 \mu\text{m}$) in order to study the circumstellar environment geometry and kinematics in the Br γ emission line. The quality of the data taken in 2007 was too poor to detect any signal in the line, whereas, as seen in Fig 2, visibility and phase variations were successfully detected in 2009. We also discovered that two other K-band emission lines were strong enough to support the kinematics study: the He I line located at $2.06 \mu\text{m}$, and the Br δ line around $1.94 \mu\text{m}$. However, the latter is located too close to the edge of the K-band so that the data SNR is significantly lower than for the two other lines. We should have decided to study the circumstellar environment geometry and kinematics

in MR mode, as it was already done for two other classical Be stars: α Ara (Meilland et al. 2007a) and κ CMa (Meilland et al. 2007b). However, since the star was bright enough, we decided to take advantage of the eight times higher spectral resolution offered by the HR mode. With this resolution of $R = 12000$, kinematics details of about 25 km s^{-1} are achievable.

Consequently, in 2010, δ Sco was observed twice with the VLTI/AMBER in HR mode. The first observations were centered on the Br γ line and the second on the He I line. Unfortunately, we did not manage to calibrate the He I observation properly. However, this does not affect either the differential visibilities and phases or the closure phase. Finally, quasi-simultaneous medium-spectral resolution observations centered on H α were carried out with the CHARA/VEGA instrument.

4.1. A qualitative analysis of the high-resolution data

The differential visibilities and phases, the closure phases (only for VLTI/AMBER data), and line profiles for these observations are presented in Fig 5. Their morphology is similar to those shown by Meilland et al. (2007a) for the Be star α Ara. The “S” shape of the phase variation and the “W” shape of some visibility variations clearly favor the hypothesis that the circumstellar environment velocity field is dominated by rotation. Using either differential visibility or phase, we can roughly determine the major-axis position angle in the plane of the sky for a purely rotating disk knowing that:

- The more pronounced “V” shape of the visibility i, the closer the baseline is to the minor axis. The more pronounced “W” shape is related to a position closer to the major-axis.
- For the same projected baseline length, the amplitude of the differential phase variation through the emission line is zero if the baseline is aligned with the minor-axis and maximal if aligned with the major-axis. This is true only if the disk is not fully resolved by the interferometer. For larger baselines, the phase amplitude first saturates and then drops, thus second order effects become visible causing the phase to lose its simple “S” shape.

Thus, both the phase and visibility variations through the observed lines favor the hypothesis of a major-axis roughly in the North-South orientation.

4.2. The kinematic model

To model the wavelength dependence of the visibility, the differential phase, and the closure phase in the observed emission lines we use a simple kinematic model developed for fast model fitting of an expanding and/or rotating thin equatorial disk. This

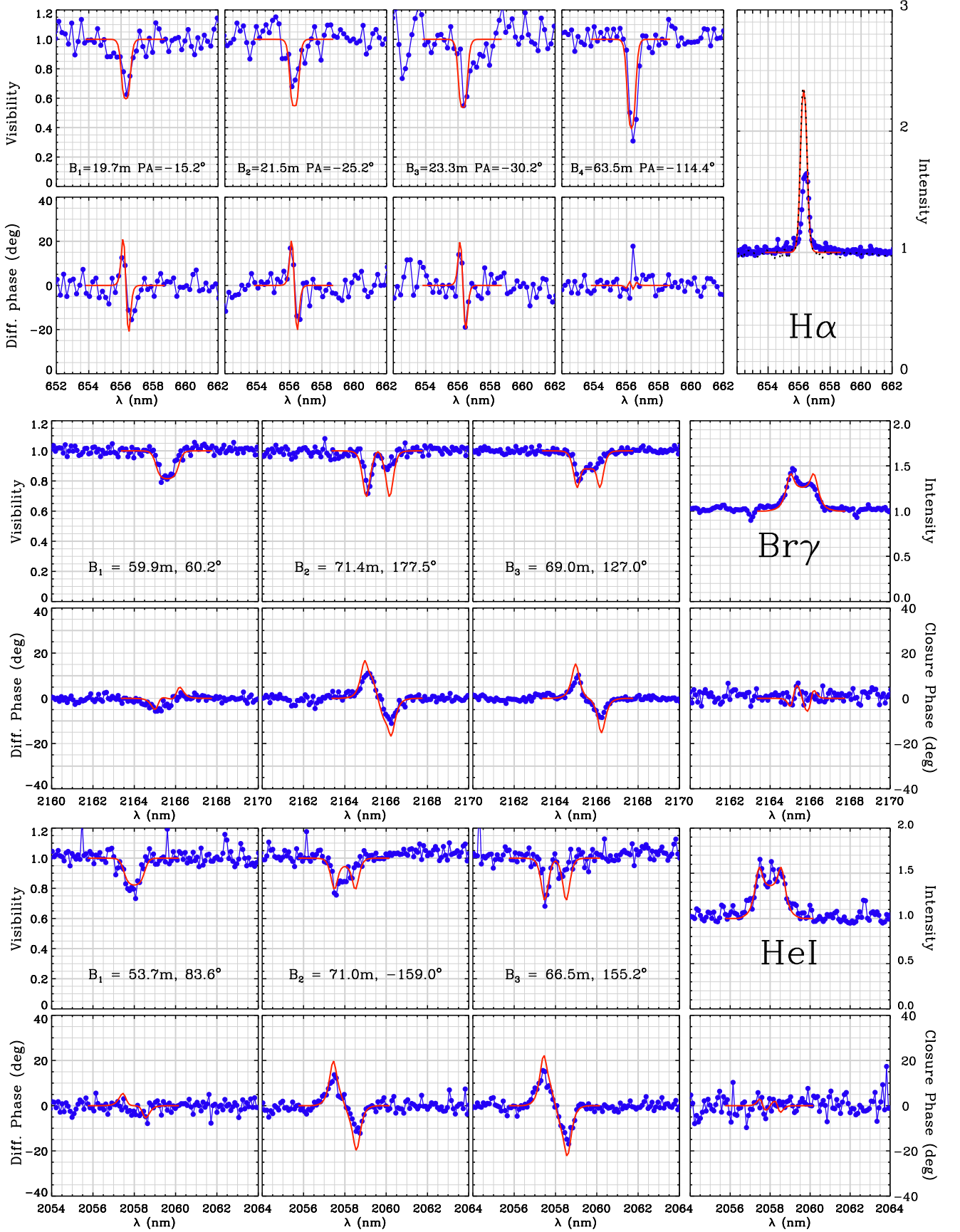


Fig. 5. Visibility and phase variations in the $H\alpha$, $Br\gamma$, and HeI emission lines obtained from the high-spectral-resolution CHARA/VEGA and VLTI/AMBER observations obtained in 2010 (blue line and circles). The best-fit kinematic-model for this dataset is over-plotted as a red line. The black dotted line (upper right figure) corresponds to an $H\alpha$ profile taken from the BeSS database and used to determine the real line EW.

model is described in detail in Delaa et al. (2011). The star is modeled as a uniform disk, the envelope emission in the continuum and the emission line has an elliptical Gaussian distribution with a flattening due to a projection effect of the geometrically thin equatorial disk, i.e., $f = 1/\cos(i)$, where i is the object inclination angle. The radial and azimuthal velocities are given by:

$$V_r = V_0 + (V_\infty - V_0) \left(1 - \frac{R_\star}{r}\right)^\gamma \quad (5)$$

$$V_\phi = V_{\text{rot}} \left(\frac{r}{R_\star}\right)^\beta \quad (6)$$

For each spectral channel in the line an iso-velocity map projected along the line of sight is then calculated and multiplied by the whole emission map in the line. Finally the whole emission map for each wavelength consists of the weighted sum of the stellar map, the disk continuum map and the emission line map within the spectral channel under consideration. The map is then rotated by the major-axis PA, and scaled using the stellar radius and distance. A 256x256x100 data-cube (i.e. 256x256 for 100 wavelengths) can be computed in less than one second on a standard computer. Finally, visibilities, differential phases, and closure phases are extracted using two-dimensional fast-Fourier transforms (FFT). We note that the pixel size was set to avoid sampling problems.

The model free-parameters can be classified into 4 categories:

1. The global geometric parameters: stellar radius (R_\star), distance (d), inclination angle (i), and disk major-axis position angle (PA).
2. The global kinematic parameters: stellar rotational velocity (V_{rot}), expansion velocity at the photosphere (V_0), terminal velocity (V_∞), exponent of the expansion velocity law (γ), and exponent of the rotational velocity law (β). These parameters describe the global disk kinematics and thus also should not depend on the observed line.
3. The disk continuum parameters: disk FWHM in the continuum (a_c), disk continuum flux normalized by the total continuum flux (F_c). These parameters depend on the observed wavelength. As the continuum visibility is highly affected by the companion and cannot be constrained properly with our few measurements we decided to neglect this contribution in our modeling. That is why we have chosen to plot differential visibilities (i.e., visibility in each spectral channel divided by the mean visibility) instead of calibrated visibility in Fig 5.
4. The disk emission line parameters: disk FWHM in the line (a_l) and line equivalent width (EW). These parameters are different for the three emission lines.

Consequently, for simultaneous fit of these three emission lines, we need a model with a total of 15 free parameters. To reduce the number of free parameters we have finally decided:

- to set the distance to that derived from Hipparcos (von Leeuwen 2007).
- to use typical values for the expansion velocity law, i.e., $\gamma = 0.86$ and $v_0 = 0$ (Stee & Araujo 1994; Stee et al. 1995).
- and to estimate the stellar radius from the fit of the Spectral Energy Distribution (SED). In Carciofi et al. (2006) they successfully fit the SED using $R_\star = 7 R_\odot$, $T_{\text{eff}} = 27000$ K, and $d = 123$ pc. In Section 3.3, we found that the distance is better given by the new Hipparcos parallax estimation from van

Leeuwen (2007), i.e. $d = 150$ pc. We managed to fit the SED using this distance, the same T_{eff} as in Carciofi et al. (2006), and a stellar radius of $R_\star = 150 \div 123 \times 7 \approx 8.5 R_\odot$.

Moreover, the lines EW are easily and efficiently constrained by the spectra plotted in Fig 5 except for the VEGA/CHARA $H\alpha$ line for which the intensity is underestimated by a factor 2-3 due to a saturation of the photon counting algorithm that is affecting the line EW but not the corresponding visibilities, as already outlined in Delaa et al. (2010). Thus the "true" $H\alpha$ line profile used to compute the EW was taken from the BeSS database⁴ where we found spectra recorded at the same epoch as our interferometric VEGA/CHARA measurements. Finally, running hundreds of models, we tried to constrain the nine remaining free parameters. Values of the best-fit model parameters are presented in Table 4. The corresponding differential visibilities and phases, closure phases, and lines profiles are overplotted in Fig 5.

Table 4. Values of the best-fit kinematic model parameters.

Parameter	Value	Remarks
Global geometric parameters		
R_\star	$8.5 R_\odot$	from the fit of the SED
d	150 pc	from von Leeuwen (2007)
i	30 deg	from the fit of the binary
PA	-12 ± 7 deg	
Global kinematic parameters		
V_{rot}	$500 \pm 50 \text{ km s}^{-1}$	$\approx V_c$
V_0	0 km s^{-1}	from Stee et al. 1995
V_∞	0 km s^{-1}	$< 10 \text{ m.s}^{-1}$
γ	0.86	from Stee et al. 1995
β	0.5 ± 0.1	Keplerian rotation
Hα disk geometry		
$a_{H\alpha}$	$9.0 \pm 3.0 R_\star$	$= 4.8 \pm 1.5 \text{ mas}$
$EW_{H\alpha}$	$7.0 \pm 1.0 \text{ \AA}$	
Bry disk geometry		
a_{Bry}	$5.5 \pm 1 R_\star$	$= 2.9 \pm 0.5 \text{ mas}$
EW_{Bry}	$6.5 \pm 0.5 \text{ \AA}$	
He I disk geometry		
a_{HeI}	$4.5 \pm 0.5 R_\star$	$= 2.4 \pm 0.3 \text{ mas}$
EW_{HeI}	$8.5 \pm 0.5 \text{ \AA}$	

4.3. The circumstellar disk extension

We managed to constrain significantly the disk extension in $H\alpha$, Bry, and He I. Modeled as a Gaussian distribution the respective FWHM are $a_{H\alpha} = 9.0 \pm 3.0 R_\star$, $a_{Bry} = 5.5 \pm 1 R_\star$, and $a_{HeI} = 4.5 \pm 0.5 R_\star$. The disk extensions in the infrared lines are compatible with those derived by Millan-Gabet et al. (2010) from 2007 Keck-Interferometer measurements, i.e. $a_{Bry} = 3.6 \pm 0.6 R_\star$ and $a_{HeI} = 4.2 \pm 0.8 R_\star$. Moreover, the size of the $H\alpha$ emission obtained from our CHARA/VEGA measurements is also roughly compatible with the Millan-Gabet et al. (2007) estimation using the line EW, i.e., $14.9 R_\star$. Consequently, it seems that the disk did not grow or shrink during the 2007-2010 period.

However, if we compare our measurements to the physical modeling by Carciofi et al. (2006) from 2005 photometric measurements, i.e. disk outer radius of $7 R_\star$, it appears that the disk might have grown between 2005 and 2007 by a factor of at least 1.3. Assuming that the extension of the $H\alpha$ emission represents

⁴ <http://basebe.obspm.fr>

the physical outer edge of the circumstellar disk, we can estimate the disk expansion velocity. Between 2000 and 2005 the disk has grown from the stellar surface (i.e. $1 R_*$) to $7 R_*$. This represents an average velocity of 0.24 km s^{-1} . Between 2005 and 2007 the disk has grown by an additional $2 R_*$, which corresponds to an expansion velocity of 0.19 km s^{-1} . These values are the same order of magnitude as those derived for the same object by Miroschnichenko et al. (2003) using the separation of double-peaked emission line profiles assuming Keplerian rotation, and for another variable Be star, Achernar (α Eri), by Kanaan et al. (2008). This value also fully agrees with the model of a viscous decretion disk proposed by Lee et al. (1991) to explain the formation of Keplerian disks around fast rotation Be stars.

Applying a 0.2 km s^{-1} expansion velocity between 2007 and 2010, we should have obtained a disk extension of $12.3 R_*$ in 2010 which is marginally larger than the extension found in $H\alpha$. However one have to keep in mind that if the disk extends too far from the central star, the density and temperature would drop too low to produce a noticeable $H\alpha$ emission.

4.4. The circumstellar disk kinematics

Using our simple kinematic model we managed to put constraints on the disk expansion and rotation velocity laws. As already stated in Section 4.1, the disk kinematics are fully dominated by rotation. An upper limit of 10 km s^{-1} can be set on the expansion velocity. This is compatible with the value of about 0.2 km s^{-1} derived in the previous section using arguments on the disk evolution.

Our estimated stellar radius, i.e. $R_* = 8.5 R_\odot$, is an average photometric radius. Thus, if the star is rotating close to the breakup velocity its equatorial radius will be larger, i.e. on the order of $10 R_\odot$, whereas the polar radius will be smaller, i.e. about $7 R_\odot$. Using the stellar mass estimated from the binary orbit, i.e. $M_* = 14 M_\odot$, the stellar breakup velocity is about $V_{\text{crit}} = 500 \text{ km s}^{-1}$. Consequently, as $V_{\text{rot}} \approx V_{\text{crit}}$ and $\beta = 0.5$, the circumstellar matter surrounding the δ Sco circumstellar disk appears to be in Keplerian rotation.

4.5. The disk asymmetry

To our knowledge, these are the first nearly simultaneously measured spectrally dispersed visibilities with the corresponding differential phases for 3 visible and infrared emission lines, namely $H\alpha$, He I ($2.06 \mu\text{m}$) and $\text{Br}\gamma$. These lines are formed at various distances from the central star as already measured for γ Cas in the $H\alpha$, $H\beta$ and He I 6678 lines by Stee et al. (1998) and as seen in the previous section where we found $\phi_{H\alpha} > \phi_{\text{He I}} > \phi_{\text{Br}\gamma}$. But compared to Stee et al. (1998) and other recent work (Carciofi et al. 2006; Millan-Gabet et al. 2010) we have now a direct access to the kinematics within each line-forming region with a velocity resolution of about 25 km s^{-1} .

We see from Fig 5 that the differential phases for these 3 lines are, first, well represented by our very simple kinematical model, and second that they are all exhibiting a typical "S" shape of a rotating disk. The amplitude of this "S" shape is on the order of $\pm 10^\circ$ which is very similar to the differential phases obtained by Carciofi et al. (2009) for the Be star ζ Tau. To the contrary of ζ Tau differential phases, our differential phases are all symmetrical. ζ Tau phases showed a very asymmetrical "S" shape with a smaller amplitude in the blue part of the "S" curve with respect

to the central line wavelength. This was a clear evidence for a one-armed spiral structure in the ζ Tau disk.

On the other hand, despite these symmetrical differential phases, our visibilities across these 3 spectral lines are asymmetrical, especially for $\text{Br}\gamma$ and He I lines, with a blue wing of the visibility systematically smaller than the red wing. The line profiles are also clearly asymmetrical, with $H\alpha$ and $\text{Br}\gamma$ line profiles showing a $V/R > 1$ whereas the S/N and spectral resolution for He I are not high enough to exhibit any line asymmetry.

Thus, δ Sco's circumstellar disk is globally rotating, with a larger (in size) and brighter emitting region in the blue part of the line, i.e. coming towards us, and a smaller (more compact), fainter or more absorbed region in the red part of the line, i.e. rotating away from us. These emitting regions are also responsible for the $H\alpha$ and $\text{Br}\gamma$ line profiles with $V/R > 1$. Moreover, since the differential phase is symmetrical, it means that the photocenter of the emitting regions are symmetric with respect to the central star (or the rotational axis). Therefore, the one-armed oscillation scenario would difficultly match the inhomogeneities detected in δ Sco's disk since it would have produced non-symmetrical spectrally resolved visibilities and phases as measured by Carciofi et al. (2009) for ζ Tau. In our case, the binarity of the system should also play an important role in the shaping and the kinematics of the circumstellar environment. For example, one could think of a tidally-warped disk from previous periastron passages of the companion star (like what is described in Moreno et al. 2011). This definitely needs a dedicated modeling effort to really constrain the detected disk inhomogeneities.

5. Conclusion

Using all available binary separation measurements including our VLTI/AMBER ones, we check the consistency of previous estimates of the δ Sco orbital elements. We found parameters that totally agrees with the latest work from Tycner et al. (2011).

Thanks to the combination of high spectral and spatial resolution we managed to constrain for the first time, not only the geometry but also the kinematics of the δ Sco circumstellar environment. As in the case of Be stars with stable or nearly stable disks such as α Ara (Meilland et al. 2007a), κ CMa (Meilland et al. 2007b), Ψ Per, and 48 Per (Delaa et al. 2011), the line emission originates from a dense equatorial disk dominated by rotation. As for most of these objects, the rotation appears to be Keplerian, with an inner boundary (photosphere/disk interface) rotating at the critical velocity (V_c). The expansion is negligible and considering an outburst scenario, should be on the order of 0.2 km s^{-1} .

Considering the measured $v \sin i$, i.e. 175 km s^{-1} , and the measured inclination angle, i.e. $30.2 \pm 0.7^\circ$, the star rotates at about 70 % of its critical velocity. This could indicate that the stellar rotation is not the main process driving the ejection of matter from the stellar surface. However, taking into account possible underestimation of the $v \sin i$ due to gravity darkening (Townsend et al. 2004), the star may rotate faster, up to $0.9 V_{\text{crit}}$.

Finally, as the measured extension in the $H\alpha$ line, $4.8 \pm 1.5 \text{ mas}$, is on the order of the binary separation at the periastron, i.e. $6.14 \pm 0.07 \text{ mas}$ (i.e., according to Tycner 2011), we expect strong interaction between the previously ejected matter and the companion.

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⁵ Available at <http://www.jmmc.fr/searchcal>

⁶ Available at <http://cdsweb.u-strasbg.fr/>

⁷ LITpro software available at <http://www.jmmc.fr/litpro>